

Now, $D(k, \omega) = 0$

$$\frac{\partial D}{\partial k} \cdot dk + \frac{\partial D}{\partial \omega} d\omega = 0$$

$$\Rightarrow \frac{d\omega}{dk} = - \frac{\partial D / \partial k}{\partial D / \partial \omega}$$

Now, $D = 1 + \frac{2\pi G \Gamma_0 |k|}{\omega^2 - k^2 - k^2 c_s^2}$

$$-\partial D / \partial k_r = \frac{-2\pi G \Sigma_0 |k| (2k c_s^2)}{(\omega^2 - k^2 - k^2 c_s^2)^2}$$

$$\partial D / \partial \omega = \frac{-2\omega (2\pi G \Gamma_0 / k)}{(\omega^2 - k^2 - k^2 c_s^2)^2}$$

$$\boxed{v_{gr, r} = -\sin(kr) (\pi G \Sigma_0 - |k| c_s^2) / (\omega - m\Omega)}$$

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$$V_{gr} \sim \underbrace{(-\sin k)}_{\oplus \text{ trailing}} \underbrace{(\pi G \Sigma_0 - |k r| C_0^2)}_{\oplus \text{ long wave}} \underbrace{(\omega - n \Omega)^{-1}}_{\substack{\downarrow \\ \oplus \text{ leading} \quad \oplus \text{ long wave} \quad r > r_{co} \\ \oplus \text{ leading} \quad \ominus \text{ short wave} \quad r < r_{co} \quad \ominus}}$$

$$k < \frac{\pi G \Sigma_0}{C_0^2}$$

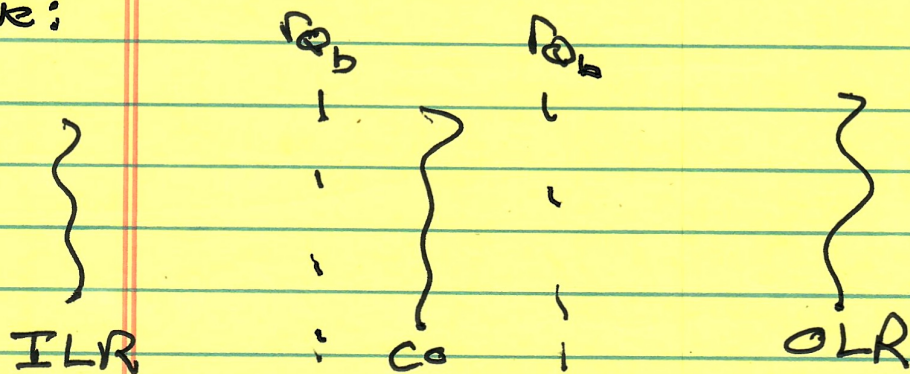
Thus:

long trailing spiral→ propagates inward for $r > r_{co}$ → propagates outward for $r < r_{co}$ long trailing spiral = always propagates
toward co-rotationshort trailing spiral→ propagates outward for $r > r_{co}$ → propagates inward for $r < r_{co}$

$\Rightarrow$ short trailing spiral always
propagates away from co-rotation.

This brings us to Amplifying Spirals:
(WATER) - over-refraction

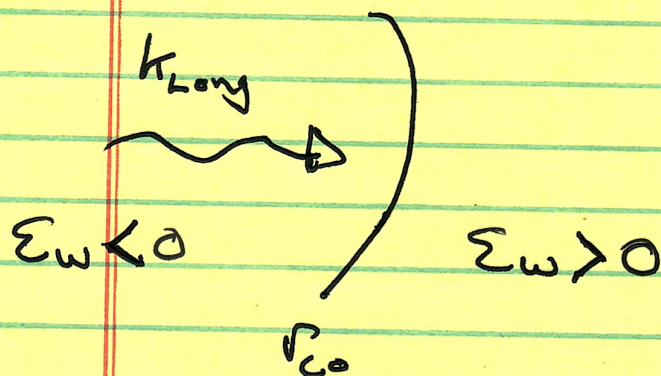
have:



Long $\rightarrow$ always propagates toward r_{co}

Short $\rightarrow$ always propagates away from r_{co} .

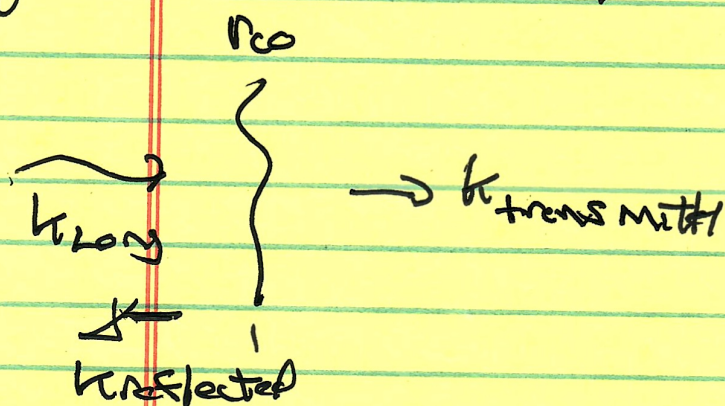
→ "Over-Refraction"



$Q \approx 1$
 → ignore Q
 barrier
 (simplicity)

Long to co.

so, also freshman physics:



so clearly reflected wave must

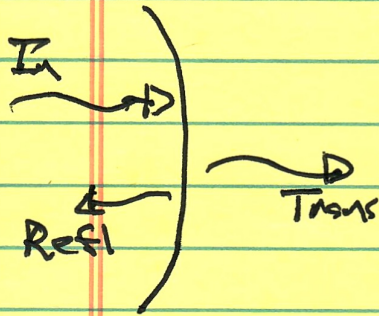
be short, so propagate away from
 co-rotation

→ Long → short conversion at
 co-rotation

→ Now, must conserve wave energy density flux across co-rotation:

$$V_{gr}^L \sum_{IN}^L = V_{gr}^S \sum_{REFL}^S + V_{gr}^S \sum_{TRANS}^S$$

$$r < r_{co} \qquad r > r_{co}$$



Now:

- if $\Sigma W > 0$, $\sum_{REFL}^S < \sum_{IN}^L$
(as in freshman physics....)

necessarily, as energy lost to transmitted wave

but → the point!:

→ if $\Sigma W < 0$. → negative energy incoming wave!

$$\Rightarrow V_{gr}^L \Sigma_{IN}^L = V_{gr}^S \Sigma_{Ref1}^S + V_{gr}^S \Sigma_{Trans}^S$$

$$\Sigma_{Ref1}^S = \left(\frac{V_{gr}^L}{V_{gr}^S} \right) \Sigma_{IN}^L - \Sigma_{Trans}^S$$

⊖

⊕

∴ Σ_{IN} is
NEW

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$$\Sigma_{Ref1}^S < 0$$

then for $V_{gr}^L \sim V_{gr}^S$ ✓

$$|\Sigma_{Ref1}^S| > |\Sigma_{IN}^L|$$

i.e. Magnitude of wave energy must
decrease!

de $E_0 = \left(\frac{\partial D}{\partial \omega} \right) \omega |A|^2$

(A) must increase!

→ Life is more interesting than in freshman physics!

→ reflected signal amplified!

→ Negative energy wave crucial!

→ NR:

- Energy source is differential rotation

- Not exponential growth, as in linear instability

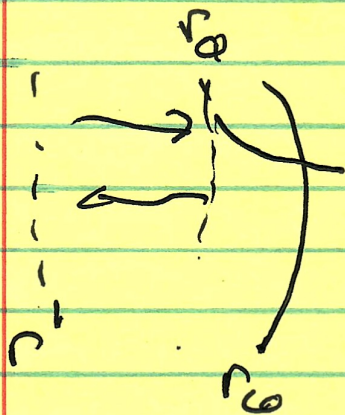
Now what of Q barrier?

- if tunneling, then weak amplification at each scattering

- but if can reflect from

$r < r_{co}$, can build up standing

wave $\rightarrow$ WASER oscillator



$$r_{ILR} < n' < r_{co}$$

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- with Q barrier:

$$\Sigma_0 \quad \Sigma_R \quad \Sigma_T$$

$$-1 = -2 + 1$$

before

now:

$$-1 = -(1+G) + \epsilon$$

↓

tunneling
reduced

[over reflected]

$G =$ tunneling
factor

$$\epsilon \sim \exp \left[-2 \int_{Q_-}^{r_0} dr |k(r)| \right]$$

→ Need find an inner reflection point:

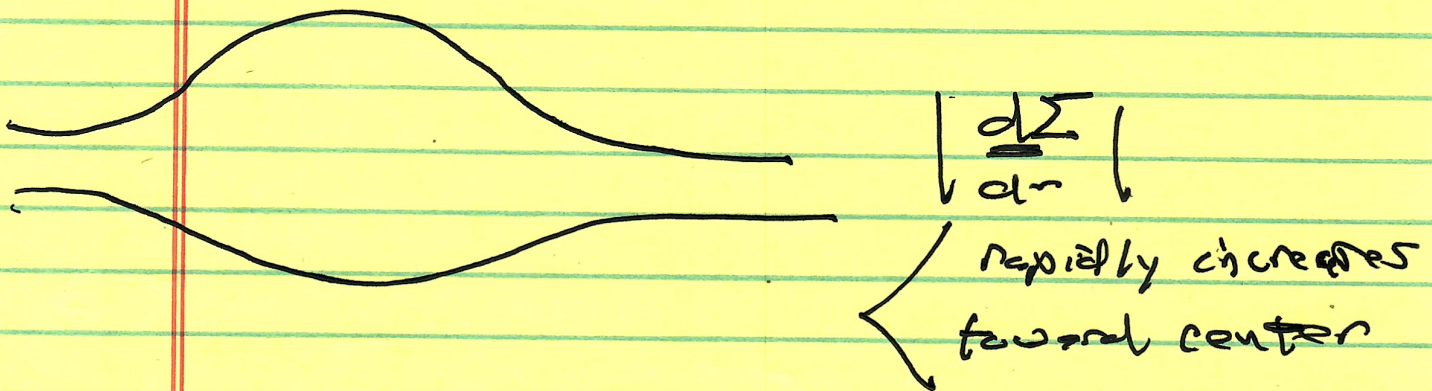
= recall:

$$V_g = -\text{sgn}(k) (\pi G \Sigma_0 - k r C_s^2) / \omega$$

$$k r > \frac{\pi G \Sigma_0}{C_s^2} \equiv k_r \rightarrow \text{short wave}$$

acoustic term dominates for short wave.

so if bulge:



$$V_g \rightarrow 0 \text{ when } \pi G \Sigma(r) = k r C_s^2$$

→ defines r_{IRP}
 ↑
 inner reflection point

where $\sqrt{\epsilon_L} < \sqrt{\epsilon_{IP}} < \sqrt{\epsilon_-}$

at $\sqrt{\epsilon_{IP}}$, $\xrightarrow{\text{inward}}$

propagating short wave converts to outward propagating long wave.

$\therefore$ restarts cycle!

50 WASER feedback loop $\leftrightarrow$ standing wave/ amplification mechanism.

i.) long wave incident on

$$\sqrt{\epsilon_-} < r < \sqrt{\epsilon_0} \quad \text{from } r < \sqrt{\epsilon_0}$$

Wave is NEW

ii.) over-reflected short wave

back-scattered to $r < \sqrt{\epsilon_0}$, propagating away from $\sqrt{\epsilon_0}$, amplified $|A|^2/A_0^2 \sim 1 + \epsilon$

(ii.) back scattered short wave comes to ν_{IRP} , due to bulge. Short wave backscattered (no transmission) as long wave propagating toward ν_{co} .

(iii.) return to (i.)

$\Rightarrow$ WAsER $\leftrightarrow$ [Negative energy wave is key element.]

N.B.

- much not understood despite long story

- amplification, NL saturation

$\Rightarrow \sum$ profile dynamics.

heating

Next: - Lindblad resonances and
spiral interaction there.

→ Goldreich + papers

- Angular momentum exchange
at Lindblad resonances

→ Spiral kinetics: Lynden-Bell
and Holmeys.

All good things must come to an
end

→ Look to proto-planetary disks
as venue for future work
on spirals.

1) Basic Scales, Ideas

Viscous flow in Cylindrical Geometry
Viscous Stress

2) Viscous Stress Model of Accretion in Thin Disk

- a) Disk structure
- b) Velocity and time scales
- c) Surface density and angular momentum transport
- d) $\dot{M}$, ν , Σ relations, alpha model
- e) Viscous Heating and Luminosity

3) Dynamics of Accretion

- a) Relaxation
 - Fixed mass $\rightarrow$ solid body rotation
 - Accretion $\rightarrow$ mass to center + 1 particle at large radius end state
- b) Rayleigh
 - 2 particles
 - Interchange at constant angular momentum, each
 - Radial buoyancy
 - Rayleigh Discriminant/Epicyclic Frequency
- c) L-B + P
 - 2 particles $\rightarrow$ conserve **sum** of angular momentum, mass
 - $\Delta E < 0 \rightarrow$ accretion + angular momentum transport outward $\rightarrow$ end state
- d) “Donkeys”, ala’ Lynden-Bell
- e) Elastically coupled particles $\rightarrow$ to MRI

4) Crash Course in MHD

- a) Equations, especially induction
- b) Alfven Theorem, Freezing-in
- c) Stresses, Energetics
- d) Waves, especially Shear Alfven
- e) Magnetic Braking, Torsional Alfven Wave
- f) Virial Theorem for MHD
- g) Partially ionized MHD + Ambipolar Diffusion
- h) Energy Principle (Introduction)

5) Magnetorotational Instability, Turbulence, Alpha, Saturation

- a) Basic Physics of MRI, Toy Model of Basic MRI, Connection to Accretion
- b) Detailed Linear Theory MRI — resistivity, partial ionization, viscosity
- c) MRI-Induced Transport:
 - Intro to Mixing Length Theory, Application to MRI

- Alpha scalings and underlying physics issues
- d) Open Issues in MRI Turbulence
- e) Magnetic Buoyancy in Disks, Parker Instability
- f) Magnetized Disk Coronae and Layered Disks, Nanoflares

6) Planetesimal Formation

- a) Drag Mechanisms, Epstein-Stokes drag, dust particulate evolution
- b) Sedimentation in Turbulence — Fokker-Planck → Subdisk thickness
- c) Goldreich-Ward Model, including Toomre stability criterion for subdisk
- d) Radiational infall
- e) Outlook on planetesimal formation

7) Basics of Galactic Dynamics

- a) Scales, Phenomena → Collisionless Stellar Dynamics
- b) Vlasov-Poisson system
- c) Physics of Vlasov Equation
- d) Jean Theorems — Weak and Strong
- e) Stationary Solutions, mostly spheroidal
- f) Jeans Equations, Virial Theorem
- g) Landau Damping
- h) Violent Relaxation and its foundations in phase mixing

8) Spiral Density Waves

- a) Motivation, Lin-Shu Hypothesis
- b) Basic Analysis
- c) Resonances: Lindblad and Co-Rotation
- d) Wave Action and Energy-Momentum Theorems
- e) Amplification Mechanisms
- f) Spirals and Angular Momentum Transport — return of L-B's "2 particles"
- g) Spirals in Vlasov Theory
- h) Outlook on spiral structure